

Searching out and studying sources of low-frequency radio emission near infrared objects

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Received May 26, 2000; accepted October 20, 2000.

Abstract. A procedure of sampling objects by way of cross-identification of the low-frequency radio catalogue of objects obtained with the aid of the Texas interferometer at the frequency 365 MHz and of the IRAS catalogue is discussed. Statistical properties of different subsamples of sources are investigated. From the results of cross-identification with the radio catalogues of the database CATS, continuous radio spectra of objects are plotted, identification with optical and X-ray catalogues is carried out. From the steep spectrum objects sources are selected for further studies.

Key words: radio sources: general, radio sources — infrared emission

1. Introduction

An important route to understanding physical nature of cosmic objects is opened up by investigations in new ranges of the electromagnetic spectrum which were previously inaccessible to astronomers. The astrophysical catalogues of the sources in IR, UV and X-ray ranges are now used frequently for searching out coordinate coincidence with radio sources and optical objects. The sieving of radio catalogues through other catalogues — cross-identification, may be helpful in detecting unusual objects. Using the CATS database (Verkhodanov et al., 1997; Trushkin et al., 1997a,b) which includes over 150 catalogues, and specially developed procedures of search, the authors performed cross-identification (Trushkin and Verkhodanov, 1996) of three catalogues of IR sources (PSC, FSC and SSC) obtained in the IRAS all-sky survey (Beichman et al., 1988) and of a radio catalogue obtained in the Texas sky survey at 365 MHz (Douglas et al., 1996).

The three principal IRAS catalogues used in this paper in the database CATS are the catalogue of point sources PSC with 245889 objects, the catalogue of faint sources FSC with 235935 objects and the serendipitous catalogue SSC obtained in the specified fields with 43886 objects. The catalogues contain data of the all-sky surveys in four ranges (12, 25, 60 and 100 μm) and also identification with the catalogues of stars, quasars, galaxies, infrared and radio sources.

The Texas catalogue of discrete radio sources (Douglas et al., 1996) brighter than 250 mJy at the frequency 365 MHz was obtained in the range $-35^\circ < \text{Dec} < +71.5^\circ$. It involves accurate posi-

tions, flux densities and parameters of the structure of 66841 sources.

The radio properties of faint objects of the IRAS catalogue (FSC) had already been investigated by way of cross-identification with the objects of the Green Bank Atlas (Condon et al., 1995) at the frequency 4.86 GHz, selected according to the criterion $S(60\mu\text{m}) \geq S(12\mu\text{m})$. As a result a catalogue of 354 sources has been compiled and a statistically significant number of radiobright FIR galaxies and quasars has been detected. Besides, in the VLA observations at the frequency 1.425 GHz with different telescope configurations Condon et al. (1996) obtained an atlas of IRAS bright galaxies.

The aim of the authors of this paper was to study objects both galactic and extragalactic near the IRAS sources, which have first of all a low-frequency radiation. This would help separate infrared objects with steep radio spectra, AGN and active galactic objects. The presence of a great number of various radio-range catalogues in the CATS database would be useful in examining in more detail the properties of the sample obtained.

2. Cross-identification. Catalogue

The cross-identification procedure consisted in searching for a match in the IRAS catalogues in the areas of different radius for each source of the Texas catalogue. We specified several radii of identification ($R=3, 10, 20, 30, 40, 60$ and $120''$) in order to follow the growth of the number of objects falling within the given circle. The coordinate errors of the sources were

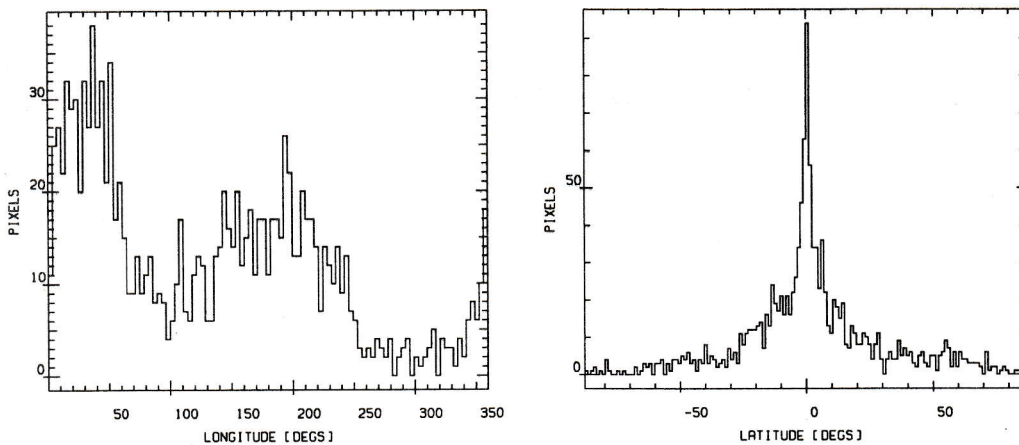


Figure 1: The distribution of sources against galactic latitude (right panel) and longitude (left panel) for the PSC sample from the ITA catalogue.

disregarded in cross-identification, although they are known to be varied from $40''$ to $80''$ for IRAS, while in the Texas catalogue with a coordinate accuracy of $\sim 1''$, shifts of the positions up to $52''$ are possible due to the ambiguity in choosing the interferometer lobe. A complete sample of the ITA (IRAS-TEXAS association) catalogue was primarily obtained for a window of $R=120''$, whereas complete radio astronomy statistical studies were performed for the catalogue obtained with a window of $R=60''$.

Simple estimates of the probability of fall-in within the circle of radius $120''$ for 133000 sources of the PSC in the region with $|b| > 10^\circ$ yield $P = 0.012$, while for the remaining 113000 sources in the region with $|b| > 10^\circ$ $P = 0.063$. Nevertheless, we have to use a window of radius over $1-2'$ for identification of extended sources in the galactic plane. The probability of random identification in cross-identification still remains. As a result, the ITA catalogue has been compiled which comprises 884 sources from the FSC, 1263 sources from the PSC and 13 ones from the SSC (208, 626 and 13 of them, respectively, have been identified for the first time). The ITA catalogue underwent the procedure of cross-identification with the radio astronomical, X-ray and optical catalogues of the CATS database, which resulted in the construction of radio spectra and investigation of statistical distinguishing features of the ITA catalogue sample. In Fig.1 is presented the distribution of objects of the PSC sample from the ITA catalogue versus galactic latitude and longitude (b, l).

The resulting catalogue of identifications of radio sources with the PSC, FSC and SSC obtained with the window of radius $60''$ contains 715 objects and includes all objects falling within the indicated zone with allowance made for coordinate errors, which increases to a certain degree the separation of the centres of gravity.

Continuous spectra of 715 objects are plotted in this paper, taking account of all radio catalogues available in the CATS database. For statistical studies the authors have selected 495 objects (Table 1) for which the distance between the centres of gravity of the Texas catalogue objects and the associated IRAS objects is less than $60''$. The rest 220 objects picked up in the original sample were not considered since they were selected taking into account the coordinate determination errors. The information about these objects is also available in the CATS database.

40 of the objects found have a single point in radio identifications (at 365 MHz). 185 objects are available simultaneously in the PSC and FSC catalogues. Besides, among the random identifications the authors are likely to have found 4 stars of class A (2 A0, A3, A5), 3 stars of class B (B0 and 2 B9), 2 stars of class F (F0 and F5), 3 stars of class G (2 G5 and G9), 51 stars of class K (from K0 to K7), 6 of them having a single radio point, 2 stars of class M (M0 and M5).

In Table 1 is presented a catalogue of radio sources for which the distance from the centre of gravity of the infrared source does not exceed $60''$. The columns indicate the following: 1 — the name of the source in the Texas catalogue, 2 — the name of the source in the IRAS catalogue, 3 — the coordinates in the Texas catalogue for the epoch 1950.0, 4 — the seconds of the right ascension and minutes and seconds of the declination of coordinates in the IRAS catalogue for the epoch 1950.0, 5 — the distance between the centroids of the sources in the Texas and IRAS catalogues in seconds of arc, 6 — the galactic latitude in degrees, 7 — the galactic longitude in degrees, 8 — the flux density at the frequency 365 MHz in mJy, 9, 10, 11 and 12 — the flux density at the wavelengths 12μ , 25μ , 60μ and 100μ , respectively, 13 — the spectral index at 365 MHz calculated as the inclination of the tangent to the radio spectrum at the given point, 14

— the spectral index at 1400 MHz, 15 — the formula of approximation of the spectrum, 16 — the possible identification or the type of object.

3. Radio properties of objects

The identification of objects of the ITA catalogue obtained by way of cross-identification of the data of the IRAS and TEXAS catalogues (Trushkin & Verkhodanov, 1996; 1997; Verkhodanov, Trushkin, 1997), made with the use of the CATS database (Verkhodanov et al., 1997), made it possible to investigate both the statistical characteristic properties of the given sample and the radio spectra of each source. The cross-identification with all radio catalogues included also the cleaning of spectra from inaccurate, in the authors' opinion, flux measurements and spectrum approximation using the standard set of curves, which enabled accurate enough evaluation of spectral indices and fluxes at the specified frequencies.

The following functions were used for approximation of spectra:

- $y = A + Bx$,
- $y = A + Bx + Cx^2$,
- $y = A + Bx + C \times \exp(x)$,
- $y = A + Bx + C \times \exp(-x)$,

where $x = \log \nu$, $y = \log S$, ν is the frequency (MHz), S is the flux density (Jy). The type of approximation was chosen automatically by the method χ^2 , after which the records were interactively examined in the visual mode of operation of the system *spg* (Verkhodanov, 1997). The weight of the spectrum points of the object was taken in inverse proportion to the squares of relative flux density errors.

Spectral indices of sources at different frequencies were computed as the inclination of the tangent at the given point of the approximation curve. It is interesting that with the general likeness of the distribution of spectral indices of the subsamples of point and faint sources of IRAS (PSC and FSC, respectively) from the ITA catalogue at the frequencies 365 and 1400 MHz there is a difference at 5000 MHz. For the sample of FSC a second peak at $\alpha \sim 0$ is clearly seen, and for the samples at all frequencies there is a peak near $\alpha = 0.8$ (Fig. 2). This peak may be due to the appearance of a population of extragalactic objects with flat spectra which may be identified with faint FSC sources, for instance, with unknown AGN.

The majority of the ITA catalogue sources have radio flux densities less than 5 Jy at the frequency of the Texas survey, whereas the fluxes in the IR range are varied from tens to a few thousands of Jy.

Among the sources of the ITA catalogue, for which identifications have been obtained for the first time, there exist at least two classes of "peculiar" objects: these are sources with inverse ($\alpha > 0$) and anomalously steep ($\alpha < -2$) spectra.

When identifying infrared sources with galactic ones, we used the following criterion: $\alpha(25\mu\text{m}, 60\mu\text{m}) \geq +1.5$ and $|b| > 10^\circ$ (Condon and Broderick, 1991). 138 sources from the PSC sample of the ITA catalogue satisfy this criterion. This subsample has $R = S_{IRAS}/S_{radio} < 200$ (where S_{IRAS} is the flux at decimetre wavelengths), and the mean flux density is $S_{365} = 0.6 \pm 0.6$ Jy, $S_{60} = 31$ Jy. Only 91 ITA sources from the PSC sample were earlier identified with galaxies.

3.1. The catalogue of radio sources identified with infrared sources to an accuracy of 10 arcseconds

The sample of objects in which the separation of centroids is less than 10 seconds of arc is of separate interest. This sample has the highest percentage of correct identifications.

20 of 98 objects of this sample lie in the Galactic plane ($|b| < 10^\circ$), 37 objects have X-ray radiation; one of these lies in the Galactic plane. Only one object of this sample has a single identification in the Texas catalogue. For 33 objects no identifications have been found in the known optical catalogues. 2 stars of class K have also fallen within the region of identifications of 3 arcseconds.

The known active galactic nuclei (of type BL Lac, QSO, Seyfert 1, 2), many of which have X-ray radiation, make up a large percentage of the identified extragalactic objects.

In Fig. 3 are displayed the spectra of the objects of the subsample of the sources with identification inside the window of radius $10''$, which have non-power spectra.

3.2. Galactic objects

Strong IR sources in the Galactic plane have been established to have the flux ratio $R > 500$. These sources are identified with Galactic HII regions, whereas non-thermal Galactic sources, supernova remnants, have $R < 100$ (Broadbent et al., 1989). The PSC sample of the ITA catalogue has 45 sources with $R > 500$. They are all identified with HII regions. For the ITA sources with $S_{60} > 100$ Jy the authors have checked the relationship between S_{365} and S_{60} . A significant correlation coefficient (0.6) has been obtained only for the sources with $S_{60} > 500$ Jy.

The thermal IR and radio emission in HII regions are proportional to the rates of ionization and recombination. For an HII region, which becomes opti-

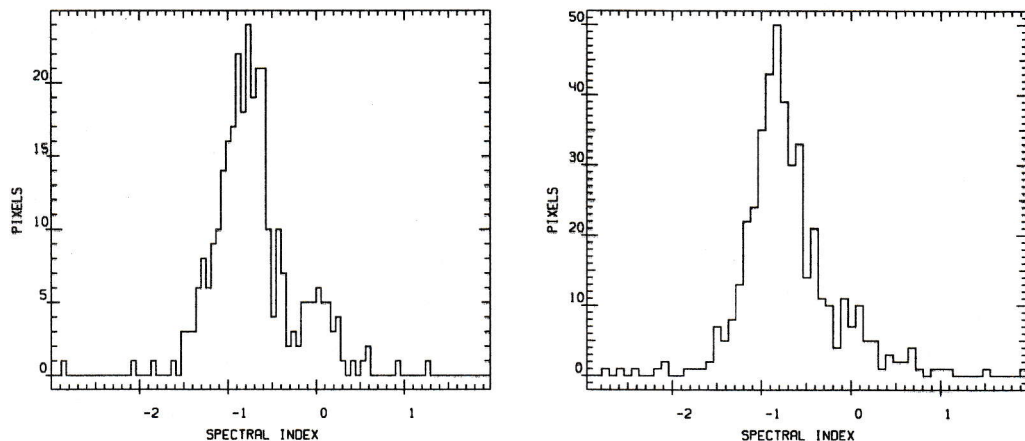


Figure 2: The distribution of spectral indices for the FSC (left) and PSC (right) samples from the ITA catalogue at a frequency of 5 GHz.

cally thin at the frequency ν , the emission capacity $\epsilon_\nu(\text{radio})$ is equal (Broadbent et al., 1989) to

$$\epsilon_\nu(\text{radio}) = 2 \times 10^{-36} T_e^{0.45} \nu^{-0.1} r W \text{cm}^{-3} \text{Hz}^{-1} \text{sr}^{-1},$$

where T_e is the electron temperature. The ratio of emissivities for $T_e=7000\text{K}$ at $60\mu\text{m}$ and at 2.7GHz is then $\epsilon(60\mu\text{m})/\epsilon(2.7\text{GHz}) = 190f$, where f is the infrared excess which is determined as the ratio of the total IR luminosity to the luminosity in L_α and, as has been shown, equal to 3–4 for typical HII regions (Broadbent et al., 1989). This ratio of emissivities at 365 MHz remains the same for an optically thin HII region, but it increases for an optically thick one.

3.3. Objects with inverse spectra

Among the objects obtained by cross-identification we have separated 31 objects (see Table 2) with spectral indices ≥ 0.1 at frequencies 360 and 1400 MHz, 18 of which lie in the Galactic plane ($|b| < 10$). 10 extragalactic sources (77 percent of the sources located outside the Milky Way) are identified with active galactic nuclei, 7 of them have X-ray radiation. The Galactic objects are likely to be identified with HII regions having a thermal spectrum.

3.4. Steep spectrum objects

In Table 3 is given a list of 130 objects having steep spectra, ($\alpha < -0.9$), at the frequencies 365 and 1400 MHz. 37 of these objects are in the Galaxy plane. For a number of the list objects a steep spectrum is frequently obtained only from measurement of fluxes in two points: at the frequencies 365 MHz (TEXAS) and 1450 MHz (NVSS). For this reason such spectra need to be confirmed by radio observations.

It should be noted that all presented objects are missing in the compiled lists of steep spectrum objects of Rötgering et al. (1994) and Chambers et al. (1996).

Eight objects have anomalously steep spectra ($\alpha < -2$): B0031-057, B0132+519, B0525+156, B1340+391, B1709+243, B1723-282, B1749-240, B1802-247. One of them, B1340+391, is observed in the survey of Bozayan (1992), however it has not been identified. In the IRAS survey in a circle of radius $60''$ is located a star of class K0 with coordinates $\alpha = 13^h 40^m 03.2, \delta = +39^\circ 12' 18$. Three objects with ultrasteep spectra are located in the Galactic plane.

4. Identification with high sensitive VLA surveys

Identification with high sensitive VLA surveys with a resolution of $5''$ (FIRST with a resolution of 1 mJy) and $45''$ (NVSS, up to 2.4 mJy) at the 21 cm wavelength has been performed. Nearly for all ITA objects (450 of 495) NVSS images are available. 44 of them are multi-component in NVSS. In the region of the FIRST survey there are observed 77 sources 52 of which are multicomponent. The multicomponent structure of objects in Galactic sources is observed in extended HII regions, and in extragalactic objects it is associated with hot spots and knots in jets.

A detailed study of the structure of objects on the basis of the data of these surveys is now under way and will be available in a separate paper.

5. Optical identifications

Cross-identification of data of different wavelength ranges is a powerful tool in searching out peculiar objects. The ITA catalogue objects were identified with both the ROSAT survey X-ray sources and the sources of the catalogue of extragalactic objects.

165 identifications with the objects of the optical catalogues AGN, PGC, MCG, the catalogue of nonstellar objects by Dixon et al. (1985) and other

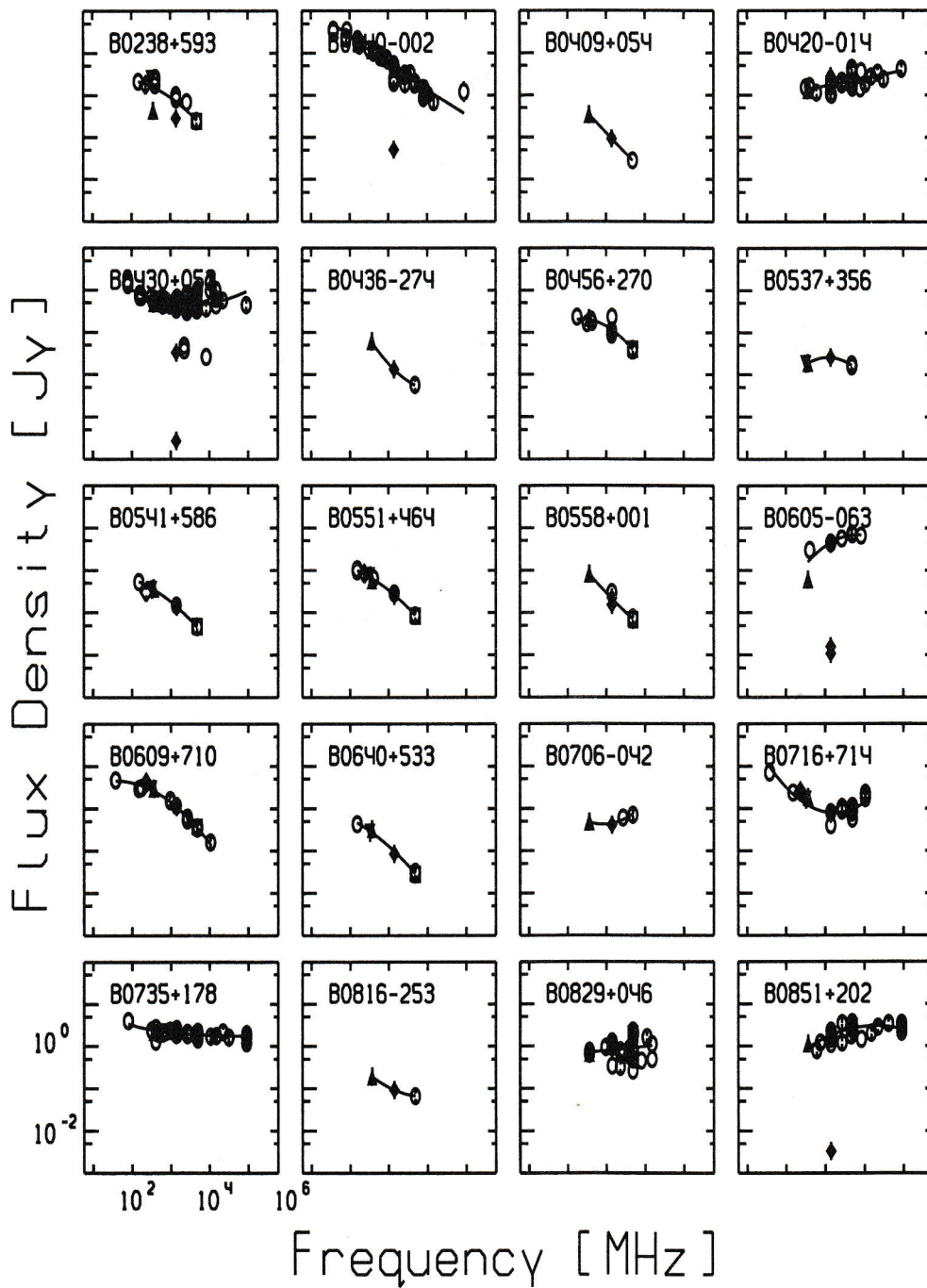


Figure 3: Radio spectra of some objects from a subsample of the ITA catalogue with identification within $10'$. The filled triangles show the points of the Texas catalogue, the diamonds are for the points of the NVSS and FIRST surveys.

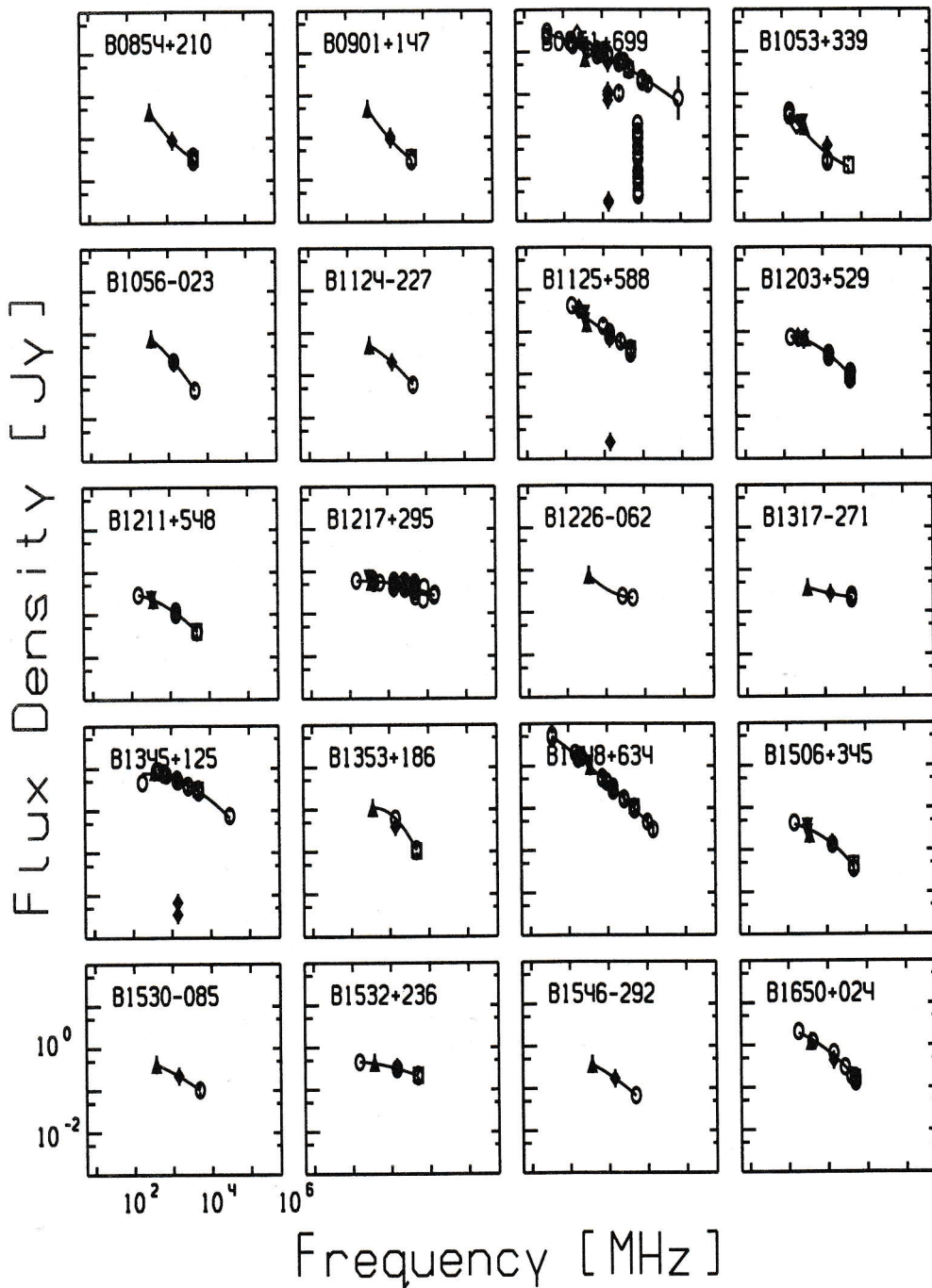


Figure 3: (Continued)

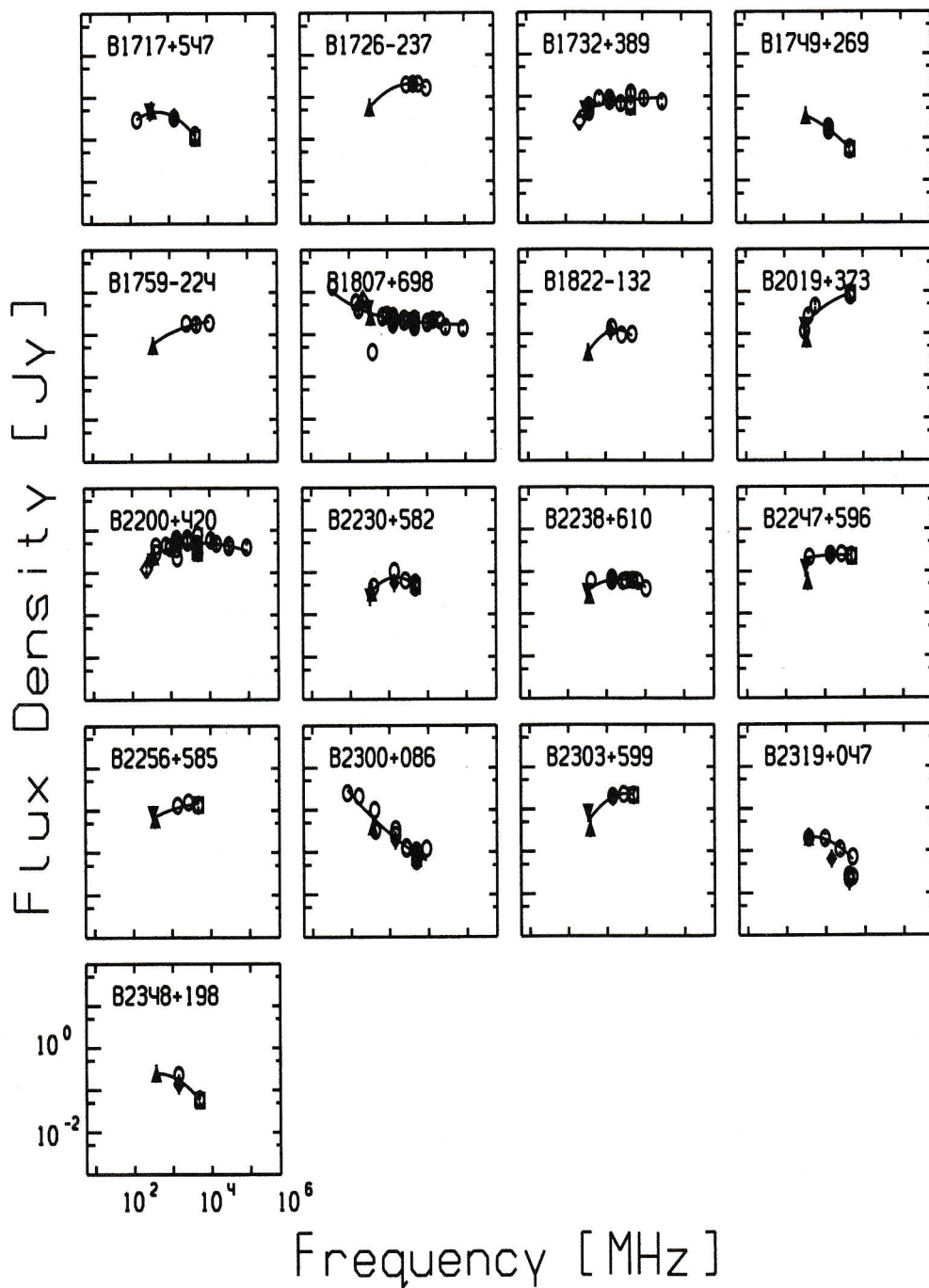


Figure 3: (Continued)

catalogues have been found. The results of the identifications are presented in Table 4.

Apart from the CATS catalogues an identification has been carried out with the Palomar Sky Survey objects with the aid of the system APM. Finding identification charts have been constructed, which are used for optical observations.

In the first column of the given table are presented the names of the sources of the Texas catalogue, the second column gives the names of the objects for which information is available in the CATS optical catalogues. In brackets are indicated stellar magnitudes if available. The reference publications are compilations in which the information about the filters used to obtain stellar magnitudes is missing. However, they can be recognized with the use of the references in the appropriate collections of papers. It was not of our concern to give full information about the presence of optical counterparts, but we just performed a cross-identification with the catalogues available in the CATS database. Attention should be drawn to 3 points: 1) the presence of a source in the list does not at all mean the existence of an optical candidate for identification, but may give a lower stellar magnitude limit if an empty field was observed; 2) not all the lists of optical objects, including those identified with radio sources, are present in CATS; this is why the data given are only the information about the objects stored in CATS; 3) part of our objects are identified with the sources of the 3C catalogue, for all 3C sources optical identifications are presently available.

In the second column of Table 4 are given the names of optical counterparts taken from the following catalogues:

3C (Spinrad et al., 1985),
 MKN (Markarian et al., 1989),
 MCG (Kogashvili, 1982),
 PGC (Paturel et al., 1989),
 QSO1 (Hewitt & Burbidge, 1991),
 QSO2 (Hewitt & Burbidge, 1993),
 UGC (Cotton et al., 1999),
 VV83 (Veron-Cetty & Veron, 1983),
 PNE SO – planetary nebulae
 (Acker et al., 1998),
 HII S – HII region catalogue (Sharpless, 1959).

For a number of sources with negative declinations there are optical data in the survey of redshifts in the southern sky (da Costa et al., 1998).

6. X-ray identifications

In identification with a circular window of radius $60''$ 82 objects of the ITA catalogue are X-ray sources, 18 of them being in the region $|b| < 10^\circ$ (see Table 5).

Tables 1 and 5 show that nearly all the X-ray sources of our catalogue (with the exception of Crab

are identified with AGN.

7. New data on the ITA catalogue objects

In the fall of 1999 new investigations of steep spectrum objects identified with the IRAS sources within $10''$ were started. Observations of 17 objects of the subsample: B0031-057, B0148+223, B0204+099, B0235+072, B0243+128, B0356+121, B0519-054, B0607+023, B0649-303, B0713+247, B1418-308, B1651-098, B1938+187, B2033-047, B2123-292, B2144-137, B2338+030 were accomplished with the radio telescope RATAN-600.

The observations were performed in 1999 November at six wavelengths during 12 days concurrently. Almost at the same time observations of spectra of 9 sources of the catalogue B0031-057, B0117+248, B0204+099, B0441+298, B2009+144, B2130-074, B2144-137, B2152-289, B2332+399, for which identifications are missing, were made under the supervision of V.H. Chavushyan with the 2.1 m telescope at the Institute of Astrophysics, Optics and Electronics (INAOE) in Mexico. The results of processing of the observations will be available in the nearest future.

One of the interesting objects of the sample, B0204+099 (IRAS FO2044+0956), is the counterpart of interacting galaxies (Fig. 4). From the results of the 2.1 m telescope observations at INAOE a redshift of 0.1 has been obtained for this pair of galaxies (Chavushyan et al., 2001).

8. Conclusions

From the results of cross-identification of the catalogues obtained at different ranges it is managed to isolate extragalactic active objects of the type of AGN. These objects may be seen in the X-ray range as well. The given approach to compiling samples may also be helpful in searching for distant dusty objects, for instance, such as IRAS F1024+4724 having a redshift $z=2.3$. This source is a distant radio galaxy and has a steep radio spectrum, and at the same time it is observed as a powerful source of infrared radiation due to dust and molecular gas. Such objects are extremely important for studying the regions of gas fragmentation into protostars. The selection of objects with the aid of the described procedure and their further investigation may increase the number of such objects.

Among Galactic objects the technique discussed isolates to advantage HII regions and planetary nebulae having low frequency radiation.

Acknowledgements. The authors are debts of gratitude to V. N. Chernenkov (SAO) and H. Andernach (Univ.

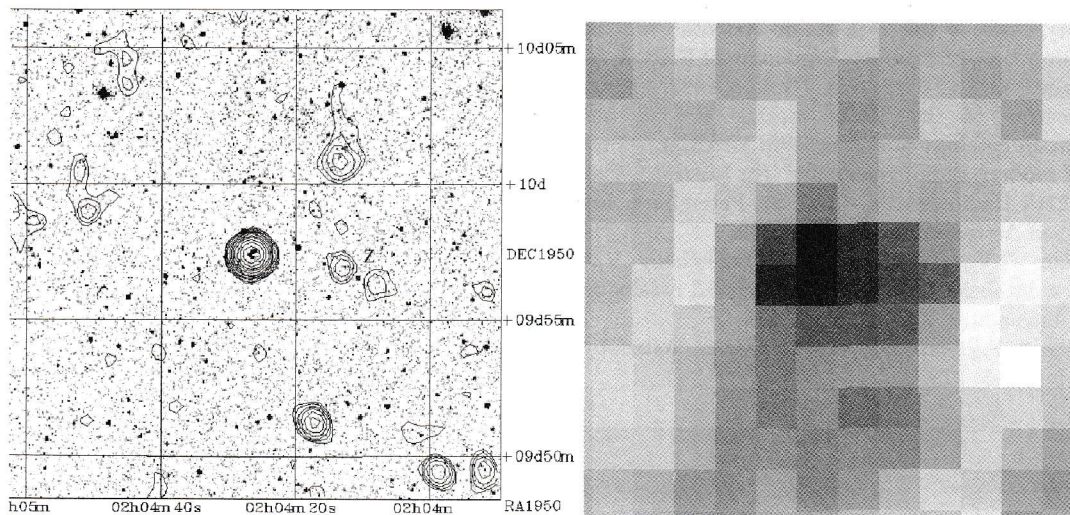


Figure 4: Upper panel: radio isophotes overlaid on the DSS $20' \times 20'$ image centered on the object IRAS F02044+0957; lower panel: IRAS grey-scale image at $60\mu\text{m}$ of the same area (pixel = $90''$). Radio isophotes are plotted in linear scale step of 0.8 mJy/beam from the level of 0.8 mJy/beam , rms of the map = 0.45 mJy/beam .

de Gunajuato, Mexico) for co-operation in establishing the CATS database and V. H. Chavushyan, R. Mujica and J. Valdes (INAOE) for making observations of the ITA catalogue objects with the 2.1 m telescope of INAOE.

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Table 2: *ITA objects having inverse spectrum*

TXS name	Flux mJy	α_{365}	α_{1400}	RA(1950) h m s	Dec(1950) o / "	IRAS name	b o	Identification
B0235+164	1031	0.48	0.31	02 35 52.57	16 24 04.3	F02358+1623	-39.109	AO 0235+164, QSO, X
B0248+430	736	0.20	0.20	02 48 18.50	43 02 56.1	F02483+4302	-14.400	AN 0248+430, QSO
B0420-014	1281	0.18	0.18	04 20 43.57	-01 27 28.3	F04207-0127	-33.139	PKS 0420-01, QSO
B0605-063	602	1.03	0.55	06 05 19.69	-06 22 31.7	F06053-0622	-12.604	X
B0641-010	215	1.07	0.41	06 41 15.56	-01 05 16.5	06412-0105	-2.218	MIN.47 06
B0829+046	671	0.10	0.10	08 29 10.89	04 39 48.4	F08291+0439	24.331	PKS 0829+046, QSO
B0851+202	1134	0.60	0.40	08 51 57.27	20 17 58.9	F08519+2017	35.821	OJ 287, QSO, X
B1404+286	179	2.02	1.11	14 04 45.89	28 41 28.4	F14047+2841	73.250	MKN 668, X
B1639-062	748	0.18	0.18	16 39 21.39	-06 15 42.7	16393-0614	25.013	
B1726-237	562	1.07	0.48	17 26 17.52	-23 43 17.5	17262-2343	5.846	PK002+05.1
B1732+389	468	0.42	0.22	17 32 40.38	38 59 47.2	F17326+3859	31.008	B2 1732+38A, X
B1759-224	546	0.88	0.46	17 59 11.70	-22 27 53.7	17591-2228	0.059	
B1822-132	356	1.38	0.31	18 22 52.85	-13 12 01.3	18228-1312	-0.389	
B1835-055	243	0.69	0.69	18 35 36.27	-05 32 27.1	18355-0532	0.411	
B1840-036	398	0.65	0.65	18 40 38.97	-03 38 40.7	18406-0338	0.173	
B1841-043	374	0.35	0.35	18 41 43.74	-04 20 49.8	18416-0420	-0.389	
B1844-025	770	0.39	0.39	18 44 23.89	-02 31 11.4	18443-0231	-0.138	
B1850+011	924	2.36	0.99	18 50 46.35	01 10 13.3	18507+0110	0.145	
B1859+041	500	0.65	0.65	18 59 23.78	04 08 22.5	18593+0408	-0.397	
B1907+090	1140	0.86	0.86	19 07 52.87	09 01 12.1	19078+0901	-0.000	
B1911+108	801	1.43	1.43	19 11 09.09	10 48 04.7	19111+1048	0.119	
B1922-224	384	0.23	0.11	19 22 40.29	-22 25 32.9	F19226-2225	-17.199	
B2019+373	731	2.90	0.65	20 19 46.25	37 21 31.8	20197+3721	0.401	
B2023+335	1121	0.26	0.37	20 23 12.99	33 33 11.2	20232+3332	-2.367	
B2025+372	999	0.30	0.45	20 25 34.92	37 12 56.8	20255+3712	-0.623	
B2101-115	217	1.13	0.46	21 01 27.52	-11 33 47.7	F21014-1133	-34.571	PKS 2101-115, X
B2200+420	2399	0.42	0.23	22 00 39.35	42 02 08.8	F22006+4202	-10.441	VRO 42.22.01, QSO, X
B2247+596	570	0.24	0.38	22 47 30.24	59 39 04.4	22475+5939	0.580	HILM
B2256+585	619	0.89	0.28	22 56 36.14	58 30 56.8	22566+5830	-0.949	HILM
B2303+599	349	1.87	0.64	23 03 04.71	59 58 26.8	23030+5958	0.045	HILM
B2323+422	283	0.46	0.24	23 23 29.30	42 15 36.2	F23234+4215	-17.600	PK 106-17.1

Table 3: *ITA-objects having ultra steep spectrum (continued)*

1	2	3	4	5	6	7	8	9
	mJy			<i>h m s</i>	<i>o ' "</i>		<i>o</i>	
B1044-170	1213	-2.00	-1.19	10 44 23.129	-17 02 26.07	F10443-1701	36.2678	
B1100-316	437	-1.27	-1.27	11 00 57.910	-31 41 32.6	F11009-3141	25.548	
B1221+343	314	-1.15	-0.98	12 21 49.945	+34 21 45.76	F12217+3422	80.9248	TXSO 1221+343
B1301+195	360	-1.01	-1.01	13 01 19.336	+19 32 22.35	F13014+1932	81.6483	MSLO ARC 1668
B1315-155	179	-1.05	-1.05	13 15 14.773	-15 33 02.78	13153-1532	46.5802	
B1331+068	1816	-1.05	-1.05	13 31 48.301	+06 49 51.27	F13317+0649	67.0607	
B1340+391	277	-3.15	-3.15	13 40 08.867	+39 11 54.83	F13400+3912	74.1143	TXSO 1340+391
B1418-308	3458	-1.21	-1.21	14 18 58.676	-30 50 46.93	14189-3050	27.9224	
B1433-050	2344	-1.05	-1.05	14 33 11.660	-05 03 09.38	14331-0502	48.9384	
B1443+107	266	-1.04	-1.04	14 43 06.977	+10 44 43.37	S144234+1016	58.4449	
B1459+136	1253	-1.36	-1.36	14 59 15.988	+13 36 36.68	F14592+1336	56.6681	
B1526-307	340	-2.82	-2.82	15 26 53.453	-30 43 01.57	15269-3042	20.7387	
B1549+076	246	-1.01	-1.01	15 49 37.227	+07 40 19.52	F15496+0740	42.9933	
B1603-235	837	-1.20	-1.08	16 03 58.301	-23 30 41.14	F16038-2330	20.6873	
B1651-098	346	-1.15	-1.15	16 51 39.648	-09 48 31.32	F16516-0948	20.4992	PGC 59239
B1655-058	406	-1.01	-1.01	16 55 10.844	-05 53 02.49	16551-0552	21.9383	
B1705-337	768	-1.36	-1.36	17 05 39.926	-33 45 59.25	17054-3346	3.7744	
B1709-257	1173	-1.30	-1.05	17 09 54.816	-25 47 41.15	17098-2547	7.7398	
B1709+243	1126	-2.30	-1.06	17 09 07.621	+24 18 33.37	F17091+2418	32.0717	
B1711-052	421	-1.22	-1.22	17 11 00	-05 17 53.8	17109-0518	18.9112	
B1715-183	582	-1.25	-1.25	17 15 13.285	-18 19 54.38	17153-1819	10.9855	
B1721+135	467	-1.13	-1.13	17 21 37.113	+13 30 33.87	F17216+1330	25.3866	
B1723-282	454	-2.44	-2.44	17 23 55.773	-28 12 57.16	17239-2812	3.7949	
B1727-338	958	-1.19	0.57	17 27 55.488	-33 51 02.71	17279-3350	-0.0526	
B1734-292	507	-1.63	-1.63	17 34 34.008	-29 16 09.53	17346-2915	1.2681	
B1737-287	386	-1.12	-1.12	17 37 31.285	-28 46 42.49	17374-2847	0.9871	
B1739+162	1190	-1.03	-1.03	17 39 11.797	+16 16 41.99	F17391+1617	22.6169	
B1749-240	382	-2.62	-2.62	17 49 36.898	-24 00 19.49	17497-2400	1.1795	
B1750-353	582	-1.03	-1.03	17 50 52.879	-35 21 15.45	17507-3520	-4.8544	
B1801-217	843	-0.31	-1.22	18 01 37.465	-21 47 24.14	18015-2146	-0.0932	
B1802-247	771	-2.17	-2.17	18 02 06.289	-24 46 05.02	18021-2445	-1.6615	
B1810+066	358	-1.00	-1.00	18 10 25.008	+06 41 15.39	18104+0641	11.6154	
B1811+498	201	-1.13	-1.13	18 11 35.465	+49 48 15.69	F18115+4948	26.4	
B1813-263	398	-1.56	-1.56	18 13 00.977	-26 21 19.9	18130-2621	-4.5832	
B1813+217	927	-1.13	-1.13	18 13 08.484	+21 46 56.7	18130+2146	17.4119	
B1817-326	389	-1.23	-1.23	18 17 12	-32 36 45.94	18171-3236	-8.3139	
B1820-049	391	-1.13	-1.13	18 20 52.789	-04 59 22.55	18207-0459	3.9118	
B1822-090	681	-1.39	-1.39	18 22 09.039	-09 04 36.39	18221-0905	1.7118	
B1827+093	830	-1.11	-1.11	18 27 51.227	+09 19 54.12	18279+0920	8.9364	
B1831+130	601	-1.13	-1.13	18 31 12.383	+13 00 27.29	18312+1300	9.8270	
B1840-156	1650	-1.08	-0.92	18 40 36.734	-15 37 36.28	18406-1537	-5.3177	
B1843-030	6780	-1.51	-1.51	18 43 46.203	-03 03 05.89	18437-0302	-0.2446	
B1844+392	322	-1.27	-1.27	18 44 32.758	+39 12 04.8	F18445+3912	17.7123	
B1855+089	292	-1.10	-1.10	18 55 43.695	+08 55 19.12	18557+0855	2.6151	
B1900+143	705	-0.90	-1.25	19 00 02.969	+14 19 03.24	19001+1418	4.1488	
B1907+077	1338	-1.12	-1.12	19 07 52.977	+07 43 34.43	19077+0743	-0.6035	
B1908+374	586	-1.17	-1.02	19 08 57.914	+37 27 53.23	19089+3728	12.6278	
B1922+100	555	-1.06	-0.94	19 22 34.344	+10 04 00.71	19226+1004	-2.7026	
B1929+377	443	-1.07	-1.07	19 29 51.359	+37 43 24.95	19299+3743	9.0357	
B1938+187	423	-1.04	-1.04	19 38 31.953	+18 44 02.44	19384+1843	-1.8623	
B1943-077	449	-1.18	-1.18	19 43 25.727	-07 42 32.28	19432-0742	-15.5878	
B2021+177	420	-1.03	-1.08	20 21 26.922	+17 45 53.57	F20213+1746	-11.0384	
B2026+062	328	-1.19	-1.19	20 26 49.852	+06 14 41.8	F20268+0614	-18.3809	
B2027+145	1154	-1.08	-1.08	20 27 38.352	+14 32 40.06	F20275+1432	-14.0626	
B2033-047	586	-1.07	-1.07	20 33 45.477	-04 47 43.74	F20337-0447	-25.3997	PGC 65054, MCG-01-52-010
B2040+082	527	-1.07	-0.94	20 40 23.758	+08 14 26.84	F20404+0813	-20.1508	
B2115+197	583	-0.95	-1.02	21 15 31.578	+19 47 21.21	21154+1947	-20.0379	
B2123-292	2972	-1.10	-1.10	21 23 43.039	-29 14 08.23	F21237-2913	-45.1483	
B2144-137	385	-1.14	-1.14	21 44 21.844	-13 46 06.75	F21443-1346	-44.976	
B2338+030	2563	-1.00	-1.00	23 38 56.898	+03 00 48.21	F23389+0300	-55.2153	PKS 2338+03

Table 4: Result of search for optical counterparts within a radius of 10' in CATS optical catalogs

name	optical counterparts
B0004+080	PGC 516 (15.4), PGC 525 (15.3), MCG+01-01-035, MCG+01-01-036
B0005-262	VV83 PKS 0005-262 (20)
B0045-255	NGC 253 (7.34), PGC 2789
B0046+316	NGC 262, MCG+05-03-008, MKN 348, PGC 2855 (14.3), QSO1
B0055+300	NGC 313, NGC 315, NGC 316, PGC 3455 (12.2)
B0104+321	NGC 383, PGC 3981 (14.6), VV83 3C 31, PGC 3982
B0105-177	PGC 4007 (14.5)
B0121+035	NGC 520, PGC 5193 (12.2), MCG+01-04-052, QSO2
B0131-296	NGC 613, PGC 5849 (10.7)
B0134+329	VV83 3C 48 (16.2), QSO2
B0140+133	NGC 660, PGC 6318 (11.7)
B0148+223	NGC 695, PGC 6844 (13.9)
B0149+359	MCG+06-05-029, MCG+06-05-031, PGC 6957 (14.3)
B0200-127	VV83 PKS 0200-127 (18.5)
B0207+389	NGC 828, MCG+06-05-092, PGC 8283 (13.1)
B0226-038	VV8Q PKS 0226-038 (17.5)
B0235+164	VV83 AO 0235+164 (19), QSO2
B0238+593	PGC 10217 (14.5)
B0238-084	NGC 1052, PGC 10175 (11.4)
B0240-002	M77, NGC 1068, VV83 3C 71, PGC 10266 (9.7)
B0241+622	VV83 4U 0241+61 (16.7), QSO1
B0243+128	PGC 10486 (14.8)
B0244-304	NGC 1097, PGC 10488 (10.3), VV83 PKS 0244-304 (10.6)
B0247+414	PGC 10792 (13.5)
B0248+430	VV83 0248+43 (15.5)
B0252-003	PGC 11007 (14.1), NGC 1144, MCG+00-08-047, PGC 11012 (13.8), QSO1
B0256+366	MCG+06-07-027, MKN 1066, PGC 11341 (13.7), QSO1
B0300+470	VV83 (18), QSO2, VV8BL 4C 47.08
B0313+411	VV83 (14.3), PGC 12171 (14)
B0313-026	PGC 12131 (13.5)
B0316+413	MKN 1505 (12.4), VV83 3C 84 (13.2), NGC 1275, PGC 12429 (12.7), QSO1
B0326-288	VV83 PKS 0326-288 (17.5)
B0338-214	VV83 PKS 0338-214 (18)
B0339-186	MCG-03-10-037, PGC 13589 (14.8)
B0345+699	MCG+12-04-006, PGC 13943 (15.1)
B0352-206	NGC 1482, PGC 14084 (13.1)
B0404+035	VV83 3C 105
B0409+054	MCG+01-11-013, PGC 14651 (15.4)
B0418+236	VV83 (20)
B0420-014	VV83 PKS 0420-01 (18), QSO2

Table 4: Result of search for optical counterparts within a radius of 10' in CATS optical catalogs (continued)

name	optical counterparts
B0430+052	MKN 1506 (14.3), PGC 15504 (15), VV83 3C 120, QSO1
B0456+270	VV83 B2 0456+27 (18)
B0502-103	VV83 PKS 0502-10 (15.4), QSO1
B0531+219	3C 144
B0537+356	HILM, PNESO 173.7+02.7
B0541+586	UCGC 03351
B0549-074	NGC 2110, PGC 18030 (13.2), QSO1
B0551+464	PGC 18078 (14.9), UGC 03374 (13.9), QSO1
B0557+653	MCG+11-08-008, PGC 18312 (14.3)
B0609+710	MKN 3 (13.8), PGC 18722 (13.9), QSO1
B0648+275	VV83 B2 0648+27 (14.9)
B0649-303	PGC 19728 (15)
B0656+330	VV83 B2 0656+33 (16.5)
B0716+714	VV83 (11), QSO2
B0717+194	VV83 OTL 0717+195 (18)
B0722+300	VV83 B2 0722+30 (15.6), QSO1
B0722-095	NGC 2377, PGC 20948 (13.5), VV83 3C 178, QSO1
B0735+178	VV83 PKS 0735+17 (16.5), QSO2
B0746-261	PGC 21863 (15.3)
B0754+100	VV83 PKS 0754+100 (16.5), QSO2
B0801+052	MKN 1210, MCG+01-21-009, PGC 22641 (14.4)
B0806-103	VV83 3C 195 (18.8)
B0816-253	PGC 23303 (11.8)
B0825-202	VV83 PKS 0825-20 (18)
B0829+046	VV83 PKS 0829+046 (18), QSO2
B0836+299	PGC 24369 (15.5), VV83 B2 0836+29B (14), QSO1
B0840+503	NGC 2639, MCG+08-16-024, PGC 24506 (12.6), QSO1
B0840+184	MCG+03-22-024, PGC 24485 (15.2)
B0851+202	VV83 (14.5), QSO2 (20)
B0901+147	MKN 1224, PGC 25476 (14.6)
B0910+403	NGC 2782, PGC 26034 (12.3)
B0943-140	NGC 2992, PGC 27982 (13.1), QSO1
B0944+045	VV83 (16.5)
B0947+248	VV83 B2 0947+24
B0951+699	M82, NGC 3034, PGC 28655 (9.2), VV83 3C 231, QSO2
B0958+559	NGC 3079, PGC 29050 (11.5)
B0958+290	VV83 3C 234 (17.5), QSO1
B1001-062	MCG-01-26-014, PGC 29192 (13.5)
B1030+161	MCG+03-27-060, PGC 31159 (14.4)
B1056-023	LCRS B1056-0219 (16.97)
B1100+282	NGC 3504, MCG+05-26-039, PGC 33371 (11.7), VV83 B2 1100+28 (11.8)
B1117+138	NGC 3628 (9), PGC 34697 (10.4),

Table 4: Result of search for optical counterparts within a radius of $10''$ in CATS optical catalogs (continued)

name	optical counterparts
B1125+588	MCG+02-29-020 (V=11.5) NGC 3690, MKN 171, PGC 35321 (12), PGC 35325, PGC 35326 (12.1), MCG+10-17-002 (V=16), MCG+10-17-003 (V=11.8)
B1127+005	VV83 PKS 1127+005 (20.6)
B1149+487	NGC 3932, MCG+08-22-023, PGC 37194 (15)
B1203+529	NGC 4102, MCG+09-20-094 (V=11.8), PGC 38392
B1211+548	NGC 4194, MKN 201, PGC 39068 (13)
B1216+507	VV83 4CP 50.33A (16)
B1217+295	NGC 4278, PGC 39764 (11.1), VV83 B2 1217+29
B1223+129	NGC 4388, MCG+02-32-041, PGC 40581 (11.8)
B1226+023	VV83 3C 273 (13), PGC 41121, QSO2
B1231+024	NGC 4536, PGC 41823 (11.1)
B1237+224	VV83 PKS 1237+224 (1)
B1244-255	VV83 PKS 1244-255 (18), QSO2
B1253-055	VV83 3C 279 (16.8), QSO2
B1254+571	MKN 231, PGC 44117 (14.3), QSO1
B1305-241	PGC 45566, QSO1, VV83 PKS 1306-241 (16),
B1308+326	VV83 B2 1308+32 (19), QSO2
B1308+373	NGC 5005, MCG+06-29-052, PGC 45749 (10.6), VV83 B2 1308+37
B1309+210	VV83 PKS 1309+21 (15.5)
B1317-271	NGC 5078, MCG-04-32-001 (V=12.5), PGC 46490 (11.7)
B1317-168	PGC 46510 (14.8), PGC 46511 (14.5)
B1319-164	MCG-03-34-063, QSO1, PGC 46710 (14.5)
B1322-295	MCG-05-32-013, NGC 5135, PGC 46974 (12.8), QSO1
B1327-206	PGC 47430 (14.6), QSO1, VV83 PKS 1327-206 (16.9)
B1334-296	M83, NGC 5236, PGC 48082 (8), MCG-05-32-050 (V=8.4), VV83 PKS 1334-29
B1336+485	NGC 5256, MCG+08-25-031 (V=14.1), MKN 266, PGC 48192 (13.9), VV83, QSO1
B1342+561	MCG+09-23-004 (V=15), MKN 273, PGC 48711 (15.1), QSO1
B1345+125	PGC 48898, QSO1, VV83 PKS 1345+12 (17)
B1350+316	PGC 49258 (15.3), VV83 3C 293 (15.6, 12.2)
B1353+186	MCG+03-36-005 (V=14.8), MKN 463, PGC 49538 (14.7), QSO1,

Table 4: Result of search for optical counterparts within a radius of $10''$ in CATS optical catalogs (continued)

name	optical counterparts
B1404+286	VV83 PKS 1353+18 (16)
B1410-029	MKN 668, PGC 50352 (16), VV83 NGC 5506, MCG+00-36-028 (V=13.6), MKN 1376, PGC 50782 (12.8), QSO1
B1411+078	NGC 5514, MCG+01-36-023 (V=14.5), PGC 50809 (14.1)
B1411+034	PGC 50845 (15.2)
B1418+546	VV83 (14.5), QSO2, BL
B1435+304	VV83 B2 1435+30
B1448+169	PGC 52984 (15.2), VV83 (13)
B1448+634	PGC 52924 (14.5), QSO1, VV83 3C 305 (16)
B1449+164	PGC 53079, PGC 53082
B1451+037	NGC 5775, MCG+01-38-014 (V=13), PGC 53247 (12.5)
B1500+423	UGC 09671 (15.6)
B1506+345	PGC 54033 (15.6), MCG+06-33-022 (V=15.6), VV83 B2 1506+34A (15)
B1512+045	PGC 54448 (16.3)
B1515+340	VV83 B2 1515+34A (15)
B1530-085	MCG-01-40-001 (V=13.5), PGC 55410 (15.3), QSO1 1530-085 (15.7)
B1532+236	MCG+04-37-005 (V=14.4), PGC 55497 (14)
B1546-292	NGC 6000, PGC 56145 (13)
B1615-201	VV83
B1641+399	VV83 3C 345 (16.3), QSO2
B1643+022	VV83 PKS 1643+022 (16.5, 18)
B1650+024	NGC 6240, MCG+00-43-004 (V=14.7), PGC 59186 (13.8)
B1651-098	PGC 59239
B1657+590	MCG+10-24-084 (V=14.2), PGC 59352 (14.1)
B1726-237	NGC 6369 (13), PK002+05.1 (10.4), PNESO
B1732+389	VV83 B2 1732+38A (19), QSO2
B1753+183	NGC 6500, MCG+03-46-003 (V=13.4), PGC 61123 (13.1)
B1807+698	MCG+12-17-000 (V=14.4), VV83 3C 371 (14.2), QSO2
B1814+426	MCG+07-37-030 (V=15.9), MCG+07-37-031 (V=16.1), PGC 61664 (15.9)
B1833+326	PGC 62082 (15.6), QSO1, VV83 3C 382 (14.5, 15.5)
B1907+508	NGC 6764, MCG+08-35-003 (V=13.2), PGC 62806 (12.7)
B1949+023	PGC 63758, VV83 3C 403 (14.5), QSO1
B2024+154	VV83 PKS 2024+154 (19.5)

Table 4: Result of search for optical counterparts within a radius of $10''$ in CATS optical catalogs (continued)

name	optical counterparts
B2025+372	HILS 106
B2033-047	MCG-01-52-010 (V=13.5), PGC 65054 (13.5)
B2101-115	NGC 7009, PK037-34.1 (8.9), PNESO 037.7-34.5, VV83 PKS 2101-115 (8.4)
B2149-084	VV83 (14)
B2200+420	VV83 (14, 15), QSO2, BL LAC
B2204+099	NGC 7212, PGC 68065 (14.9), QSO1
B2223-052	VV83 3C 446 (16.9, 18.4), QSO2
B2230+582	HILS 138
B2238+610	NGC 7354, PK107+02.1 (12.9), PNESO 107.8+02.3
B2247+596	HILS 146
B2256+585	HILS 152
B2300+086	NGC 7469, MKN 1514, PGC 70348 (12.9), QSO1 (13.15)
B2303+599	NGC 1470, HILS 156
B2314+038	PGC 70899, QSO1, 3C 459
B2319+047	ZCG 2319+04 (15.9)
B2322+149	MCG+02-59-038, PGC 71370 (13.4)
B2323+422	NGC 7662, PK106-17.1 (9.2), PNESO 106.5-17.6
B2325+085	NGC 7674, MCG+01-59-080 (V=14), MCG+01-59-081 (V=16), MKN 533, PGC 71504 (13.9), PGC 71505 (15.4), QSO1
B2338+030	VV83 PKS 2338+03 (20.2)
B2348+198	NGC 7770,7771, PGC 72635 (14.5), PGC 72638 (13)
B2349-014	PGC 72664, QSO1, QSO2, VV83 PKS 2349-01 (16, 17.5)
B2353+299	ZWG 499.010 (15.2), PGC 72893 (15.1)

Table 5: Result of search for X-ray counterparts within a radius of $60''$ in CATS optical catalogs

name	X-ray counterparts
B0045-255	ENOBS, RXJ004733-2517.3
B0046+316	1WGA J0048.7+3157, RXJ004847+3157.2, ENOBS
B0055+300	1WGA J0057.8+3021, RXJ005748+3021.2, ENOBS
B0104+321	EXSS, EIN2S
B0106+612	ENOBS
B0121+035	1WGA J0124.5+0347, RXJ012434+0347.6, ENOBS
B0134+329	1WGA J0137.6+3309, RXJ013741+3309.6, ENOBS
B0140+133	1WGA J0143.0+1338, RXJ014302+1338.6
B0226-038	ENOBS, EIN2S
B0235+164	1WGA J0238.6+1637, RXJ023838+1637.1, ENOBS

Table 5: Result of search for X-ray counterparts within a radius of $60''$ in CATS optical catalogs (continued)

name	X-ray counterparts
B0238-084	1WGA J0241.0-0814, RXJ024104-0815.2, ENOBS
B0240-002	1WGA J0242.6-0000, RXJ024240-0000.7, EIN2S
B0241+622	ENHRI, ENOBS
B0244-304	1WGA J0246.3-3016, RXJ024619-3016.4, EIN2S
B0252-003	1WGA J0255.1-0010, RXJ025511-0010.7
B0313+411	1WGA J0316.7+4119, RXJ031641+4119.4, ENOBS
B0313-026	1WGA J0315.9-0225
B0316+413	RXJ031945+4130.7, ENOBS
B0338-214	ENOBS, EIN2S
B0420-014	ENOBS, EIN2S
B0430+052	1WGA J0433.1+0521, RXJ043311+0521.2, ENOBS
B0531+219	ENSNR Crab
B0549-074	ENHRI, EIN2S
B0551+464	EIN2S
B0605-063	1WGA J0607.8-0622, RXJ060747-0623.9
B0609+710	1WGA J0615.5+7102, RXJ061535+7102.1, EIN2S
B0716+714	1WGA J0721.8+7120, RXJ072152+7120.4, EIN2S
B0735+178	1WGA J0738.1+1742, RXJ073807+1742.4, EIN2S
B0754+100	EIN2S
B0829+046	EIN2S
B0835+645	1WGA J0840.1+6419, RXJ084006+6419.7
B0840+503	1WGA J0843.6+5012, RXJ084338+5012.1
B0851+202	1WGA J0854.8+2006, RXJ085448+2006.4, EIN2S
B0910+403	EIN2S
B0943-140	1WGA J0945.6-1419, RXJ094541-1419.6, EIN2S
B0951+699	1WGA J0955.8+6940, RXJ095550+6940.4, EIN2S
B0958+559	1WGA J1001.9+5539, RXJ100159+5541.0, EIN2S
B0958+290	ENQSO
B1100+282	EIN2S
B1117+138	1WGA J1120.2+1335, RXJ112015+1335.6, EIN2S
B1125+588	1WGA J1128.5+5833, RXJ112831+5833.7
B1217+295	1WGA J1220.1+2916, RXJ122007+2916.7
B1223+129	RXJ122546+1239.7

Table 5: Result of search for X-ray counterparts within a radius of 60'' in CATS optical catalogs (continued)

name	X-ray counterparts
B1226+023	1WGA J1229.1+0202, RXJ122906+0203.0, ENHRI
B1231+024	EIN2S
B1253-055	1WGA J1256.1-0547, RXJ125611-0547.3, EIN2S
B1254+571	RXJ125613+5652.2
B1308+326	1WGA J1310.4+3220, RXJ131028+3220.6, EIN2S
B1308+373	1WGA J1310.9+3703, RXJ131056+3703.3
B1322-295	1WGA J1325.7-2950, RXJ132543-2950.1
B1334-296	1WGA J1337.0-2952, RXJ133700-2952.0, ENHRI
B1336+485	1WGA J1338.3+4816, RXJ133818+4816.6
B1342+561	1WGA J1344.6+5553, RXJ134441+5553.0
B1345+125	1WGA J1347.5+1217, RXJ134732+1217.2
B1350+316	EIN2S
B1353+186	1WGA J1356.0+1822, RXJ135603+1822.2, EIN2S
B1404+286	1WGA J1406.9+2827
B1410-029	1WGA J1413.2-0312, RXJ141315-0312.3, EIN2S
B1418+546	RXJ141946+5423.0, EIN2S
B1448+634	EIN2S
B1530-085	EIN2S
B1532+236	1WGA J1534.9+2330

Table 5: Result of search for X-ray counterparts within a radius of 60'' in CATS optical catalogs (continued)

name	X-ray counterparts
B1632-281	1WGA J1635.8-2812, RXJ163552-2813.0
B1641+399	1WGA J1642.9+3948, RXJ164258+3948.6, ENHRI
B1650+024	1WGA J1652.9+0224, RXJ165259+0224.1
B1732+389	1WGA J1734.3+3857, RXJ173420+3857.7
B1737-080	EIN2S
B1802+025	E_sl
B1807+698	1WGA J1806.8+6949, RXJ180649+6949.5, EIN2S
B1830-210	RXJ183339-2103.7
B1833+326	ENHRI, EIN2S
B2002+335	RXJ200422+3339.0
B2023+335	1WGA J2025.1+3343, RXJ202510+3342.9
B2101-115	1WGA J2104.1-1121, RXJ210410-1121.8
B2200+420	1WGA J2202.7+4216, RXJ220243+4216.6, EIN2S
B2223-052	1WGA J2225.7-0456, RXJ222547-0456.9, EIN2S
B2300+086	1WGA J2303.2+0852, RXJ230315+0852.5, ENHRI
B2325+085	1WGA J2327.9+0846, RXJ232757+0846.7
B2348+198	1WGA J2351.4+2006, RXJ235124+2006.7, EIN2S
B2349-014	EIN2S